

A&A manuscript no.
(will be inserted by hand later)

Your thesaurus codes are:
3 (11.02.2, 13.07.2, 13.25.2)

ASTRONOMY
AND
ASTROPHYSICS

November 15, 1999

X-ray/TeV-gamma-ray observations of several strong flares of Mkn 501 during 1997 and implications

H. Krawczynski¹, P.S. Coppi², T. Maccarone², F.A. Aharonian¹,

¹ Max Planck Institut für Kernphysik, Postfach 103980, D-69029 Heidelberg, Germany

² Yale University, P.O. Box 208101, New Haven, CT 06520-8101, USA

Received 12 July 1999; accepted 27 October 1999

Abstract. During more than 6 months in 1997, the BL Lac object Mkn 501 was in an exceptionally bright state, both in the X-ray band and in the Very High Energy (VHE) band. In this paper we present a multiwavelength study of Mkn 501 during this extraordinary outburst. We describe the analysis of a data base of X-ray observations acquired with the pointed X-ray telescopes on board the *Rossi X-Ray Timing Explorer* (RXTE) during April, May, and July, 1997. We combine this data set with detailed VHE spectral information obtained from simultaneous or nearly simultaneous observations with the stereoscopic Cherenkov telescope system of HEGRA. Several strong flares could clearly be resolved in both energy bands, making it possible to perform a detailed correlation analysis of the X-ray and VHE γ -ray flux levels and spectra. We interpret the results in the framework of a Synchrotron Self Compton models and study the constraints on the model parameters. We show that the emission mechanism produces an Inverse Compton spectrum with a significant curvature in the TeV energy range and discuss the implications for estimates of the intergalactic extinction due to pair production processes of the VHE photons on the Diffuse Extragalactic Background Radiation.

Key words: BL Lacertae objects: individual: Mkn 501 – Gamma rays: observations – X-rays: galaxies

1. Introduction

The BL Lac object Mkn 501 ($z=0.034$) underwent a major outburst in X-ray and in the VHE bands during 1997. Observations with the X-ray telescopes on board the BeppoSAX and the RXTE satellites showed, compared to earlier data, a flux increase at several keV by up to one order of magnitude and a Spectral Energy Distribution (SED) peaking frequently at very high energies, namely in the energy range above ~ 25 keV (Pian et al. 1998a; Lamer &

Wagner 1999). In the VHE band Mkn 501 was the brightest known source in the sky showing dramatic flux variability from a fraction to approximately ~ 10 times the flux of the Crab Nebula (Aharonian et al. 1999a, called A99a in the following; Aharonian et al. 1999c; Bhat et al. 1997; Catanese et al. 1997; Djannati-Atai et al. 1999; Hayashida et al. 1998), with photon energies up to ~ 20 TeV (Aharonian et al. 1999b, called A99b in the following).

The high energy nonthermal continuum emission of BL Lac objects is widely believed to originate in a relativistic jet due to a population of high energy electrons, emitting synchrotron radiation at longer wavelengths and higher energy photons in Inverse Compton (IC) processes of the highest energy electrons with lower energy seed photons (see for recent reviews Coppi 1997; Sikora 1997; Ulrich et al. 1997). The origin of the IC seed photons has not yet been established. In so-called “Synchrotron Self Compton” (SSC) models the target photon population is dominated by low energy synchrotron photons (e.g. Bloom & Marscher 1993; Ghisellini et al. 1996; Mastichiadis & Kirk 1997). In “External Compton” models the seed photons originate outside the emission volume and are, e.g. radiation from the nuclear continuum scattered or reprocessed in the broad-line regions (see e.g. Sikora et al. 1994) or accretion disc photons (Dermer & Schlickeiser 1994). Simultaneous observations of the highest energy synchrotron photons in the X-ray band and of the highest energy IC photons in the VHE band make it possible to infer complementary information about the rapidly evolving population of highest energy electrons. Besides studies of large populations of similar sources, only detailed studies of the temporal and spectral characteristics of individual sources over a broad wavelength region promise to yield sufficient constraints to unambiguously identify the mechanism responsible for the observed emission. A further necessity to study in detail the temporal emission characteristics arises from the fact that the observed VHE γ -ray spectra are expected to be substantially modified by the intergalactic extinction due to pair production processes of the VHE photons with the Diffuse Extragalactic Background Radiation (DEBRA) (Nikishov 1962; Gould & Schröder 1966; Stecker et al. 1992). The temporal anal-

Send offprint requests to: Henric Krawczynski (email address: Henric.Krawczynski@mpi-hd.mpg.de)

Correspondence to: Henric Krawczynski

ysis should yield enough redundant information not only to identify the emission mechanism but also to determine the jet parameters making it possible to infer the electron spectrum from its X-ray synchrotron emission and to predict the intrinsic VHE spectrum. The comparison of the intrinsic and the observed spectra yields the intergalactic extinction and as a consequence an estimate of the DEBRA density in the relatively unconstrained $0.5 \mu\text{m}$ to $50 \mu\text{m}$ wavelength region. The observations with the HEGRA stereoscopic system of Cherenkov telescopes of 1997 showed that the Mkn 501 time-averaged VHE energy spectrum extends to energies well above 10 TeV and made it possible to sample an energy spectrum with an exponential cutoff deeply into the exponential regime (A99b). From 500 GeV to ~ 20 TeV the spectrum can be described by a power law model with an exponential cutoff: $dF/dE \propto (E/1 \text{ TeV})^{-0.9} \exp(-E/6.2 \text{ TeV})$. In this paper we will study in detail the question whether the cutoff is caused by intergalactic extinction or by the emission mechanism itself. Note that the HEGRA telescope system achieves an energy flux sensitivity νf_ν at 1 TeV of $10^{-11} \text{ erg/cm}^2 \text{ s}$ for 1 hour of observation time. Furthermore, it is possible to determine differential spectra with reasonable statistical accuracy for integration times of a few hours for sources with strong, Crab like VHE γ -ray flux levels. Such Cherenkov telescope installations together with pointed X-ray telescopes like BeppoSAX, RXTE, or ASCA make it possible to study the temporal and spectral properties of BL Lac objects on time scales of hours.

Earlier discussions of the implications of the 1997 Mkn 501 X-ray and VHE data can be found in Pian et al. (1998a), Tavecchio et al. (1998), A99a-b, Bednarek & Protheroe (1999), Hillas (1999), and Konopelko et al. (1999b). We present in this paper the analysis of a large data base of RXTE observations of Mkn 501 during 1997 and combine it with the simultaneous and nearly simultaneous VHE data from A99a,b. Based on the spectral and temporal X-ray data presented in this paper and the spectral information measured with the HEGRA telescopes we re-examine the constraints on SSC scenarios and their model parameters (as given e.g. in Tavecchio et al. 1998; Bednarek & Protheroe 1999). Furthermore, following an approach described in Coppi & Aharonian (1999) and Hillas (1999) we present SSC fits to the RXTE and BeppoSAX data which yield, together with the HEGRA data, information on the degree of intergalactic TeV Gamma-ray extinction.

The paper is structured as follows. The data sample and the analysis of the RXTE and the HEGRA data is described in Sect. 2. The RXTE results and the correlations of the X-ray and VHE γ -ray flux levels and spectra are investigated in Sect. 3. In Sect. 4 we describe possible SSC model scenarios and SSC fits to the data. In Sect. 5 we summarize the results.

!1651.F1

Fig. 1. Four RXTE 3-25 keV SEDs. The given MJD is centered on the observation window. From above to below the spectra have been measured with 15 min, 18 min, 29 min, and 15 min integration time.

2. RXTE and HEGRA observations and analysis

We report on public RXTE observations which were performed from April 3rd, 1997 to July 14th, 1997. During April and May two observations were made each night. The integration time of all pointings were between 10 and 70 minutes per pointing. The RXTE satellite carries two pointed X-ray experiments (Bradt et al. 1993), i.e. the Proportional Counter Array (PCA) sensitive in the 2 keV to 100 keV energy range with good sensitivity below 25 keV, and the High Energy X-ray Timing Experiment (HEXTE) with sensitivity in the energy range from 15 keV to 150 keV. Due to the very limited statistical information in the HEXTE data and possible problems with the absolute HEXTE flux normalization and response matrices we used only the 3-25 keV PCA data for the spectral fits. After applying the standard screening criteria, the spectra were extracted with *FTOOLS 4.1* using bright source background models, and spectral fits were performed with *XSPEC 10.0*. A constant neutral hydrogen column density of $2 \times 10^{20} \text{ cm}^{-2}$ was chosen, a value which lies between the 21 cm line HI result of $1.73 \times 10^{20} \text{ cm}^{-2}$ (Stark et al. 1992) and the ROSAT spectral absorption result of $2.87 \times 10^{20} \text{ cm}^{-2}$ (Lamer et al. 1996). Since the analysis is restricted to the energy region above 3 keV the chosen hydrogen column density has only a minor influence on the fitted spectra. The majority of measurements are satisfactorily fitted with single power law models; for days with long integration times and high count rates we need broken power law models to adequately describe the data.

The HEGRA Cherenkov telescope system (Daum et al. 1997; Konopelko et al. 1999a) is located on the Roque de los Muchachos on the Canary Island of La Palma (lat. 28.8° N , long. 17.9° W , 2200 m a.s.l.). During monitoring of Mkn 501 from March 16th to October 2nd, 1997, 110 h of high quality data was taken. The analysis tools and the estimate of the systematic errors on the differential γ -ray energy spectra are discussed in A99a and A99b. The results of fits to the differential spectra determined for 63 individual days are given in A99a; the 1997 time-averaged spectrum over the energy range from 500 GeV to 24 TeV is given in A99b.

3. The Mkn 501 X-ray and VHE γ -ray characteristics during the 1997 outburst

Fig. 1 shows four Mkn 501 SEDs as determined from the 3-25 keV RXTE PCA data. It can be clearly recognized that falling SEDs with spectral indices (α from $F_\nu \propto \nu^{-\alpha}$) larger than one, as well as increasing SEDs with spectral

!1651.F2

Fig. 2. The RXTE 10 keV energy flux levels in $10^{-12}\text{erg/cm}^2\text{s}$ (a), the RXTE 3-25 keV spectral indices (b), the HEGRA flux levels at 2 TeV in $10^{-12}\text{erg/cm}^2\text{s}$ (c), and the HEGRA 1-5 TeV spectral indices (d) as function of time. For April 16th, 1997 (MJD = 50554) the VHE activity has been estimated from the CAT measurement (shown by an asterisk). X-ray and VHE γ -ray observations taken with less than 6 h time difference are highlighted by solid symbols.

indices smaller than 1 have been observed. While the spectrum of MJD 50578 can be described satisfactorily by a pure power law model, the spectra of MJD 50545, 50554, and 50579 show evidence for spectral softening with increasing energy. For analyzing the trends in the spectral evolution we characterize the 3-25 keV spectral steepness by a single power law index. Note that during the 1997 outburst the position of the SED peak in the X-ray energy region varied substantially on a time scale of weeks: while for MJD 50545 (April 7th, Fig. 1a) the peak is found at ~ 5 keV, the RXTE data as well as the BeppoSAX data of MJD 50554 (April 16th, Fig. 1b) show a peak above 25 keV.

Fig. 2a-b show the fitted RXTE X-ray energy flux at 10 keV and the 3-25 keV spectral index as function of time. The flux varies by a factor of three with shortest exponential increase/decay times¹ of roughly one day. The spectral index varies from 0.7 to 1.1 with typical velocities of 0.01/h. We do not find any evidence for faster hardening than softening.

The HEGRA energy flux at 2 TeV and the 1-5 TeV spectral index α as function of time are shown in Fig. 2c-d. On April 16th, RXTE as well as BeppoSAX observed an extremely bright X-ray flare. Unfortunately, no HEGRA observations were possible, and the flux at 2 TeV has here been estimated from lower energy observations with the CAT Cherenkov telescope (Djannati-Atai et al. 1999). Although the VHE Mkn 501 flux showed variability by factors of up to ~ 30 over the whole observation period, and by factors of up to 10 for the days with nearly simultaneous RXTE observations, the spectral index – determined with an accuracy of typically 0.1–0.3 – is rather stable and statistically consistent with a constant value. Aharonian et al. (A99a, A99b) do not find any indication for a correlation of the absolute flux and the spectral shape. From 1 TeV to 5 TeV the statistical accuracy of the observations is rather high, namely $\simeq 0.05$ in the spectral indices of flux selected spectra. Between 500 GeV and 1 TeV and from 10 TeV to 15 TeV the statistical accuracy is rather modest and changes of the spectral index in these energy regions by as much as 0.3 can not be excluded. Remarkably, the mean 1-5 TeV spectral index of the data taken dur-

¹ For two flux levels F_1, F_2 measured at times t_1, t_2 , the exponential increase/decay time τ is defined as $\tau \equiv (t_2 - t_1)/\ln(F_2/F_1)$.

!1651.F3

Fig. 3. The 3 keV to 25 keV spectral index against the 10 keV energy flux. The April and May data is shown by solid symbols, the July data by open symbols.

!1651.F4

Fig. 4. The Discrete Correlation Function determined from the 3 keV and 25 keV RXTE fluxes with one day bins (a) and with 6 h bins (b). The calculation of the flux values is based on the power law fits to the RXTE PCA data. Positive Δt correspond to the 3 keV flux variability leading the 25 keV flux variability.

5.cm!1651.F5

Fig. 5. The correlation coefficient of the X-ray flux at energy E_X and the VHE γ -ray flux at 2 TeV as function of E_X . The X-ray fluxes at E_X have been determined from the power law fits to the 3-25 keV RXTE data. Only data pairs with a time difference between X-ray and VHE observation of less than 6 hours have been used.

ing phases of rapidly rising TeV flux differs by 0.10 ± 0.06 from the mean index of the data taken during phases of rapidly falling TeV flux, weakly indicating a harder spectrum during phases of increasing fluxes. The TeV data shows shortest exponential increase/decay times of about 15 h; some evidence has been found for TeV variability on shorter time scales of a few hours (A99a; Quinn et al. 1999).

On 25 days the time separation between the X-ray and the VHE γ -ray observations was smaller than 12 h and for 22 days it was smaller than 6 h. X-ray/VHE data pairs with time delays less than 6 h are marked in Fig. 2 by solid symbols. Since the X-ray and the VHE variability is found on time scales slower $\gtrsim 1/2$ day and the spectral indices changed typically by less than 0.01 units per hour these measurements are well suited for a meaningful correlation analysis. Note that data with a clear signature of a substantially increasing or decreasing flux are of utmost importance for extracting information about the acceleration and emission mechanisms. The X-ray/VHE observations during April, May, and June cover a considerable number (~ 5) of distinct strong flares including several phases of substantial flux increase and decrease.

The X-ray data shows a tight correlation of the absolute flux and the spectral index (Fig. 3). An obvious tendency is that higher fluxes are accompanied by harder spectra. Compared to the April and May data, relatively high X-ray fluxes during July were observed with – in comparison to the correlation shown by the April and May data – rather soft spectral indices, indicating that conditions in the source might have changed from April/May to July. A possible time lag between the 25 keV flux and the 3 keV flux has been searched for using the Discrete Correlation Function (DCF) (Edelson & Krolik 1988). The analysis on a time scale of days shows that the time lag is smaller than one day (Fig. 4a). On a time scale of hours

!1651.F6

Fig. 6. Correlation of the RXTE energy fluxes at $E_X = 3$ keV and $E_X = 25$ keV, left and right side respectively, with the 2 TeV energy fluxes. The X-ray fluxes have been determined from the power law fits to the 3-25 keV RXTE PCA data. The lines shows fits according to $\nu F_\nu(2 \text{ TeV}) \propto (\nu F_\nu(E_X))^2$. Only data pairs with a time difference between X-ray and VHE observation of less than 6 hours have been used.

!1651.F7

Fig. 7. The Discrete Correlation Function determined from the 25 keV RXTE and the 2 TeV HEGRA fluxes with one day bins (a) and 12 h bins (b). Positive Δt correspond to the 2 TeV flux variability leading the 25 keV flux variability.

the DCF favors a time lag of the 3 keV flux behind the 25 keV flux of smaller than $\simeq 15$ hours (Fig. 4b). Note that here and in the following the statistical error bars on the DCF have been computed according to the prescription of Edelson & Krolik (1988) from the deviations of the DCFs of individual data pairs from the mean DCF value of the corresponding bin, taking thereby the quality of the correlation into account. The errors on the DCF values resulting from the statistical errors of the measurements alone are much smaller.

Fig. 5 shows the correlation coefficient of the X-ray flux at energy E_X and the flux at 2 TeV as a function of E_X . Hereby the fluxes at E_X were estimated from the power law fits to the RXTE data. For X-ray energies above several keV we find an excellent correlation with a correlation coefficient of about 0.8–0.9. The statistical errors on the DCF do not allow us to decide whether the correlation of the 3 keV and the 2 TeV or the 25 keV and the 2 TeV flux levels is better. Fig. 6 shows the X/TeV-correlations for E_X equal to 3 keV and 25 keV. A quadratic relation between the keV and the TeV fluxes fits the data rather well.

A DCF analysis of the 25 keV and 2 TeV fluxes shows no time lag on a timescale of days (Fig. 7a). An analysis on hour time scale indicates that the 25 keV variations happen rather simultaneously with the TeV variations, or even lead them by several hours (Fig. 7b). Although a time lag of the IC radiation relative to the synchrotron radiation is generally expected in SSC scenarios (Coppi & Aharonian 1999) the DCF of Fig. 7b does not yet allow definitive conclusions about the existence of such a time lag in Mkn 501.

One of the most interesting questions is whether the X-ray and VHE spectral indices are correlated. The X-ray spectral index for all nights with VHE observations varies only by 0.25 which equals to a good approximation the median accuracy of the HEGRA spectral index estimates. We draw the very important conclusion that VHE spectral variations with spectral changes comparable to the ones observed by RXTE are not excluded by the HEGRA data. The statistical accuracy of the HEGRA spectral es-

timates can be improved by determining average spectra for several days. A closer look at the RXTE spectral indices however shows that the mean spectral index of a sufficient large number of days varies by not more than $\simeq 0.1$, a spectral difference which is again difficult to assess at TeV energies. In our best try we grouped together HEGRA observations according to an RXTE spectral index below and above 0.85, with mean spectral indices which differ by 0.1. The HEGRA data gives a 1-5 TeV spectral index of $1.20 \pm 0.02_{\text{stat}}$ for the data sample with harder X-ray spectra, and $1.28 \pm 0.04_{\text{stat}}$ for the data sample with softer X-ray spectra. A difference of the VHE spectral indices of the same magnitude as at X-ray energies seems likely but the statistical evidence is marginal. To summarize, the variations of the RXTE spectral indices are rather small compared to the accuracy of the HEGRA spectral estimates and the data supports rather than excludes a correlation of the X-ray/VHE spectral indices. This is consistent with recent claim of the CAT group who reported that the joint spectrum of the data of the two strongest April flares, was harder than the 1997 mean spectrum (Djannati-Atai et al. 1999).

4. Interpretation in the framework of SSC scenarios

In this section we first discuss possible SSC scenarios for explaining the 1997 Mkn 501 X-ray and VHE characteristics and discuss the constraints on the model parameters (Sect. 4.1). Eventually we show several examples of fits to multiwavelength data and compare the estimated emitted VHE spectrum with the observed one (Sect. 4.2) to infer information about the intergalactic extinction caused by the DEBRA (Sect. 4.3). Note that the quadratic dependence of the VHE (probably IC) γ -ray fluxes on the X-ray (probably synchrotron) fluxes shown in Fig. 6 favors a SSC scenario in which the production of seed photons is closely connected to the production of the nonthermal electron population.

4.1. Constraints on the model parameters

In the following a “one zone model” is used for simplicity; the synchrotron and IC radiations originate in a spherical emission volume homogeneously filled with an isotropic population of relativistic electrons which move in a magnetic field characterized by its mean value B .

In the first scenario (“scenario 1”), the emission volume is characterized by a constant Doppler factor ($\delta_j^{-1} = \Gamma(1 - \beta \cos(\theta))$, with Γ and β the bulk Lorentz factors of the emitting volume and θ the angle to the observer), a rather constant magnetic field B , and a slightly variable radius. While the spectrum of accelerated electrons does not change with time, only the variability of the rate of accelerated particles entering the emission region causes the observed X-ray and VHE γ -ray variability. The X-

ray synchrotron spectrum hardens when a large amount of freshly accelerated particles enters the emission region (i.e. during the rising stage of a flare) and subsequently steepens due to the cooling of these particles (i.e. during the decaying stage of a flare). A continuously changing X-ray spectral index is expected in this scenario if the cooling time of electrons responsible for the X-ray radiation is comparable to the flux variability time scale. We envisage that the electrons responsible for the low energy X-ray emission (below 1 keV or so) have no time to cool before they escape the emission region and as a consequence the spectra are always very hard in this energy region. By contrast, at energies above ~ 25 keV most often soft spectra are observed due to the rapid cooling of the responsible electrons. Diffusive shock acceleration predicts a power law spectrum of nonthermal particles $dN_e/dE \propto E^{-p}$ with $p \simeq 2$ which is steepened at higher energies by one unit ($p' \simeq 3$) due to synchrotron cooling. Therefore, the produced synchrotron radiation has a spectrum with $F_\nu \propto \nu^{-\alpha}$ with $\alpha = (p-1)/2 \simeq 0.5$ in the low energy region and with $\alpha' = (p'-1)/2 \simeq 1$ in the high energy region. Indeed, the Mkn 501 spectra observed with BeppoSAX during April 1997 show for all three days a spectrum F_ν with an spectral index $\alpha = 0.5$ below several keV, and $\alpha' \approx 1$ above $\simeq 25$ keV. Furthermore, observations with OSSE taken between April 9th to April 15th, 1997 showed a hard 50 keV to ~ 470 keV spectrum characterized by a spectral index α' of 1.1 (Catanese et al. 1997). In this scenario the X-ray spectral index is determined by the temporal evolution of the density of emitting particles rather than by the density itself. As a consequence, the correlation of the X-ray and VHE spectral indices should be tighter than the correlation of the absolute VHE fluxes and VHE spectral indices, in accord with the results shown in the previous section. Assuming an escape time of the low energy electrons smaller than several days, the IC seed photon density approximately follows the rate of injected particles and a faster than linear (but not necessarily quadratic) rise of the VHE fluxes for increasing X-ray fluxes is a natural consequence of the changing injection rate of accelerated particles.

The following alternative scenario (“scenario 2”) seems attractive since it automatically accounts for the < 1 keV X-ray spectrum with a spectral index $\alpha \approx 0.5$, repeatedly observed with BeppoSAX (Pian et al. 1998a,b). The basic difference to scenario 1 is a minimum energy E_1 of accelerated electrons which is responsible for the break of the synchrotron spectrum. A large magnetic field leads to rapid cooling and hence steepening of initial electron spectra down to the minimum energy E_1 . Below E_1 electrons which underwent cooling form a spectrum with the canonical spectral index $p'' = 2$, independent of the value p of the initial spectrum. Above E_1 the electron spectrum is characterized by a spectral index $p' = p + 1 \approx 3$ for $p \approx 2$. In this scenario one obtains naturally a synchrotron spectrum with $\alpha = (p'' - 1)/2 = 0.5$ from the electrons with

energies below E_1 . Furthermore, the scenario allows for a high magnetic field which assures fast cooling (relative to the flux variability time scale) of the electrons near the peak of the SED and thus a high radiation efficiency. Spectral changes in the X-ray energy range are caused in this scenario by the parameters E_1 and p varying with time.

We explored in detail a third scenario (“scenario 3”) in which the observed characteristics are caused by a time-independent electron spectrum, but where all 3 parameters (B , R , and the density of emitting particles) vary substantially. A changing magnetic field together with a rather stable (e.g. due to cooling on small time scales) electron spectrum could be the reason for large spectral changes in the X-ray band accompanied by a stable VHE spectrum. Note, however, that it is necessary in this scenario to introduce fixed relationships between the parameters characterizing the emission volume (i.e. between R and B) in order to explain the very tight X-ray/VHE correlation. Further studies in this direction are underway.

Note that an alternative scenario in which the spectral variability as well as the X-ray and VHE intensity variations are caused by varying only the relativistic Doppler factor δ_j of the emission volume can be excluded from the X-ray data alone. As a consequence of the Lorentz invariance of F_ν/ν^3 a change of the energy at which the Mkn 501 SED peaks by a factor of a would be accompanied by a change of the observed flux by $\approx a^{3+\alpha}$, where α is the spectral index at the frequency of observations ($F_\nu \propto \nu^{-\alpha}$). The spectra of Fig. 1 show a change of the SED peak position by approximately a factor of 5 but clearly without the corresponding increase in luminosity.

Deciding on one of the first three scenarios will be crucial for a final understanding of the Mkn 501 SEDs. In the following we will summarize the constraints on the model parameters. The significant VHE γ -ray variability on timescales Δt_{obs} of approximately half a day implies the well known upper limit on the radius of the emitting volume:

$$R < \delta_j c \Delta t_{\text{obs}} = 1.3 \cdot 10^{16} (\delta_j/10) (\Delta t_{\text{obs}}/12\text{h}) \text{ cm} \quad (1)$$

Given this limit, the Doppler factor is constrained by the condition of a negligible or modest absorption of high energy photons in pair production processes with lower energy synchrotron photons inside the emission volume. Extrapolating the synchrotron spectrum observed by BeppoSAX at soft X-ray energies ($\lesssim 1$ keV) according to $F_\nu \propto \nu^{-\alpha}$ with $\alpha = 0.5$ towards lower frequencies, we obtain (see e.g. Svensson 1987) an optical depth as function of the photon energy E_γ (frame of observer) of the gamma-ray:

$$\tau_{\gamma\gamma} \simeq (\Delta t_{\text{obs}}/12 \text{ h})^{-1} (E_\gamma/10 \text{ TeV})^{0.5} (\delta_j/4.8)^{-5}. \quad (2)$$

Here and in the following we assumed a Hubble constant of $60 \text{ km s}^{-1} \text{ Mpc}^{-1}$. If the (long wavelength) target photon density in the source changes as the observed X-ray

and VHE luminosity, one would expect a steepening VHE spectrum for large VHE luminosities. The non-observation of such a correlation with an accuracy of better than 0.1 in the VHE spectral index limits the optical depth at 10 TeV to values below $\sim 1/4$, resulting in a lower limit on the Doppler factor of $\delta_j \gtrsim 6.3$.

The X-ray and VHE γ -ray flux levels (Fig. 6) indicate that the electron cooling is dominated by synchrotron cooling rather than by IC cooling. For the case that the spectral evolution after flares is dominated by synchrotron cooling, Takahashi et al. (1996) and Kirk et al. (1998) discussed the possibility to constrain the magnetic field from observations of time lags between high and low energy synchrotron radiation. Due to the energy dependent cooling times of electrons with energy E_e : $t_s = [(4/3) \sigma_T c (B^2/8\pi) E_e/m_e^2 c^4]^{-1}$ with σ_T the Thomson cross section, one expects that after a flare the lower energy synchrotron radiation (from electrons of energy E_l) lags the higher energy synchrotron radiation (from electrons of energy E_h) by $\Delta t_s = t_s(E_l) - t_s(E_h)$. As a consequence, the DCF of the lower and higher energy radiation should show a shoulder reflecting this time lag. Using the delta functional approach for the energy of synchrotron photons E_S produced by electrons of energy E_e :

$$E_S \simeq 200 (\delta_j/10) (B/1 \text{ G}) (E_e/1 \text{ TeV})^2 \text{ keV}, \quad (3)$$

and a shoulder of the DCF of the 3 keV and the 25 keV radiation up to $\Delta t_s \approx 15 \text{ h}$ (Fig. 4) yields:

$$B \simeq 0.025 (\Delta t_s/15 \text{ h})^{-2/3} (\delta_j/10)^{-1/3} \text{ G}. \quad (4)$$

Since several other reasons are expected to broaden the DCF (e.g. delays due to the size of the emission region) this value of the magnetic field should be regarded as a rough lower limit.

In the reminder of this subsection we discuss which X-ray and VHE photons are produced by the same electrons and are henceforth expected to show correlated intensity variations. We find that the IC photons near and above the energy at which the gamma-ray SED peaks are mainly produced in Thomson scatterings as well as in the Thomson to Klein-Nishina (KN) transition regime. This can qualitatively be understood as follows (a more quantitative discussion is given in the next subsection, see also Zdziarski 1989). On the one hand, due to a probable seed photon energy flux rising approximately as $\nu F_\nu \propto \nu^{0.5}$ the Thomson processes yield the largest energy flux of IC photons for the highest seed photon frequencies. On the other hand, due to a seed photon number spectrum falling steeply like $dN/d\nu \propto \nu^{-1.5}$ with increasing frequency, and due to the suppression of the KN scatterings inversely proportional to the centre of momentum energy of seed photon and electron the KN production of VHE photons is strongly suppressed. Electrons of energy E_e produce therefore the largest IC power per logarithmic

bandwidth interacting with photons from the energy (all energies in the frame of the emitting volume)

$$\varepsilon_0 \approx 1/4 (m_e c^2)^2 / E_e = 1/16 \cdot (E_e/1 \text{ TeV})^{-1} \text{ eV} \quad (5)$$

producing Gamma-rays of energy (frame of observer):

$$E_{IC} \approx \delta_j \eta E_e. \quad (6)$$

Obviously, for IC scattering in the extreme KN regime it is $\eta \approx 1$; for Thomson scattering on photons of energy ε_0 one naively expects from the delta functional approximation of the IC cross section (e.g. Ginzburg & Syrovatski 1964) a value η of $1/3$. From our simulations we find that electrons of energy E_e indeed produce the maximum energy per logarithmic bandwidth at IC photon-energies given by Eq. (6) with $\eta \approx 1/3$ near the peak of the SED. Note that in our models the synchrotron self absorption cutoff is typically found between 10^{11} Hz and 10^{12} Hz (laboratory frame), well below the frequencies of the most important seed photons.

Eqs. (3) and (6) can be combined to derive the relation of the observed synchrotron and IC photon energies produced by electrons of the same energy region:

$$E_{IC}/\text{TeV} \approx \left(\frac{\delta_j/10}{B/0.05 \text{ G}} \right)^{1/2} (E_S/1 \text{ keV})^{1/2}. \quad (7)$$

It can be recognized that the “mapping” between synchrotron and IC photon energies is only a weak function of δ and B . It should be noted that Eqs. (3), (6), and (7) are only rough approximations; in reality, electrons of a certain energy produce synchrotron and IC photons with energies scattered over more than a magnitude.

4.2. SSC fits to multiwavelength data

Figs. 8 and 9 show simultaneous and nearly simultaneous RXTE, BeppoSAX, and HEGRA data taken on April 7th and April 13th, 1997 respectively together with the results from the SSC models. Note that the RXTE spectrum from April 7th is with a spectral index of 1.1 among the softest spectra observed with RXTE in 1997 while the spectrum of April 13th has a rather average spectral index of 0.83.

The SSC model is computed with the time-independent part of the code described by Coppi (1992). This time-independent part is similar to the codes described e.g. by Inoue & Takahara (1996) and Kataoka et al. (1999). For a given set of parameters (δ_j , B , R) the intrinsic VHE spectrum can be computed by deriving the spectrum of the emitting electron population from the synchrotron spectrum (Coppi & Aharonian 1999). This “inversion” of the synchrotron spectrum yields unique results for sufficiently smooth electron spectra. Unfortunately, the predicted VHE spectrum depends partially on the electron spectrum outside the energy region constrained by the X-ray observations. Given the observational evidence and the theoretical framework outlined

above, we use the following probable assumptions about the synchrotron spectrum outside the energy region covered by the RXTE observations, namely below 0.1 keV and above 25 keV. The RXTE observations, as well as the OSSE observations from April 9th to April 15th, 1997 show spectra with spectral indices of $\lesssim 1.1$. We choose accordingly a spectrum of accelerated electrons with $dN_e/dE \propto E^{-p}$ with $p = 2.2$. Electrons which did not have time to cool produce a synchrotron spectrum at energies below 0.1 keV with $F_\nu \propto \nu^{-0.6}$. Above 25 keV the cooled electrons give a spectrum with $F_\nu \propto \nu^{-1.1}$. First we assume a Doppler factor $\delta_j = 25$ and a magnetic field $B = 0.037$ G which lie well above the lower limits from Eqs. (2) and (4); later we show how results change varying these parameters in a wide margin. For this parameter values the 3-25 keV radiation is produced by roughly the same electrons as the 3-9 TeV emission (see Eq. 7) in accord with the tight X-ray/VHE correlation shown in Figs. 5 and 6.

The synchrotron spectrum probably cuts off at some energy. The BeppoSAX and the OSSE observations do not show any indications for such a cutoff up to energies of at least 150 keV. Importantly, for the used values of δ and B the location of this cut off does not influence the SSC prediction of the IC spectrum up to energies of 25 TeV as long as the cut off does not influence the synchrotron spectrum below ~ 150 keV (compare Eq. (7)). So in the following we use a rather high exponential cutoff at 10 MeV. Note that 10 MeV synchrotron photons can be produced by ~ 20 TeV electrons (Eq. (3)) which is still below the maximum energy of electrons expected for the case of shock acceleration with Bohm diffusion (e.g. Hillas 1999).

The predicted IC to synchrotron luminosity ratio depends on the assumed radius of the emission volume as $1/R^2$. Assuming negligible intergalactic extinction at 500 GeV we determine a radius of $1.5 \cdot 10^{16}$ cm and $1.1 \cdot 10^{16}$ cm for the April 7th and April 13th data, respectively. Doppler factors below $\simeq 15$ result in too large radii which conflict with the upper limit from the flux variability time scale (Eq. (1)). Note however that actually there could be some intergalactic extinction at 500 GeV; one should keep in mind that the fit predicts the shape rather than the absolute normalization of the intrinsic VHE spectrum.

Finally we turn to the SSC model spectra given in Fig. 8 and Fig. 9 (solid lines). The model calculations show that the break of the electron spectrum observed in the synchrotron radiation at X-ray energies is expected to give rise to a significant curvature of the IC spectrum in the TeV energy region. Obviously, the scenario favors a very small or alternatively an energy-independent intergalactic extinction of the radiation up to approximately 5 TeV. The observed time-averaged Mkn 501 spectrum shows additional steepening above 5 TeV which is not found in the

!1651.F8

Fig. 8. BeppoSAX (asterisks), RXTE (squares), and HEGRA (open circles) observations of April 7th, 1997 (MJD 50545) and the HEGRA 1997 time-averaged Mkn 501 spectrum scaled according to the detection rate of this day (full circles). The line shows a SSC model for a Doppler factor $\delta_j = 25$, a magnetic field $B = 0.037$, and a radius of $R = 1.5 \cdot 10^{16}$ cm. On April 7th, 1997, the HEGRA observations were performed from UTC 4:53 to 5:27. Two RXTE observations were made, the first from UTC 7:16 to 7:31 and the second from UTC 10:37 to 10:52 showing to an accuracy of 0.008 no significant change of the X-ray spectral index. The BeppoSAX observations were performed from UTC 5:30 to 16:00. The BeppoSAX LECS, MECS, and PCA data has been normalized to fit the RXTE data at energies around 5 keV. Compared to the spectrum shown in Pian et al. (1998a) we reduce the PCA normalization by 35% which eliminates the apparent discontinuity of the BeppoSAX spectra at $\simeq 15$ keV and is then consistent with the spectral shape measured from 3 keV to 25 keV with RXTE. The 2σ upper limit at $2.4 \cdot 10^{22}$ Hz has been derived from EGRET observations between April 9th and April 15th, 1997 under the assumption of a constant emission level (Catanese et al. 1997).

predicted spectrum. Possibly, this additional steepening is caused by the intergalactic extinction.

The BeppoSAX, OSSE, and RXTE observations show that the X-ray spectrum turns over from a spectral index $\alpha \approx 0.6$ below several keV to $\alpha' \approx 1.1$ above several keV. As can be seen in Figs. 8 and 9 the corresponding electron spectrum produces an IC component with a power law behavior well below 500 GeV (spectral index $\alpha_{IC} = 0.6$) and well above several TeV (spectral index $\alpha'_{IC} = 1.6$). The power law shape of a low and high energy IC component in an energy region where the KN effect is not negligible was already discussed by Tavecchio et al. (1998). The appearance of the IC power law components can be understood by noting that the relative contribution to the production of IC photons from scatterings in the Thomson regime and from scatterings in the KN regime does not depend on the energy of the responsible electrons as long as the seed photon and electron populations are described by power law spectra in the relevant energy ranges². As a consequence, the slope of the produced IC component can be derived by considering scattering processes in the Thomson regime alone and using the δ -functional approximation of the IC cross section (see also Tavecchio et al. 1998). The rate of produced IC photons is given by

$$\dot{n}_{IC}(E_{IC}) = c \int_{E_1}^{E_2} \int n_e(E_e) n_{ph}(\varepsilon) \sigma(E_e, \varepsilon, E_{IC}) d\varepsilon dE_e, \quad (8)$$

² The statement can easily be derived from the properties of the KN cross section, namely that the cross section as well as the fraction of electron-energy transferred to the IC photons only depends on the centre of momentum energy of seed photon and electron, i.e. on the product of the seed photon and electron energy in the comoving frame of the emission volume.

11651.F9

Fig. 9. RXTE (squares), and HEGRA (open circles) observations of April 13th, 1997 (MJD 50551) and the HEGRA 1997 time-averaged Mkn 501 spectrum scaled according to the detection rate of this day (full circles). The solid line shows a SSC model for a Doppler factor $\delta_j = 25$, a magnetic field $B = 0.037$, and a radius of $R = 1.1 \cdot 10^{16}$ cm. The dashed and dotted lines show the uncertainty in the shape of the predicted VHE spectrum arising from changing the value $\sqrt{\delta_j/B}$ (dashed line: $\delta_j = 100$, $B = 0.012$ G, dotted line: $\delta_j = 25$, $B = 0.12$ G). On April 13th, 1997, the HEGRA observations were performed from UTC 1:37 to 2:55. Two RXTE observations were made, the first from UTC 6:24 to 6:53 and the second from UTC 11:01 to 11:54 showing to an accuracy of 0.006 no significant change of the X-ray spectral index. The 2σ EGRET upper limit at $2.4 \cdot 10^{22}$ Hz is the same as in Fig. 8.

where n_{IC} , n_e , and n_{ph} are the numbers of IC photons, high energy electrons, and seed photons per unit volume and per energy interval. Using $dn_e/dE_e \propto E_e^{-p}$ for E_e between the energies E_1 and E_2 , $dn_{ph}/d\varepsilon \propto \varepsilon^{-(\alpha+1)}$ over the whole relevant energy region, and the rough approximation for the cross section for the production of IC photons $\sigma = \sigma_T \delta \left[E_{IC} - \frac{4}{3} \varepsilon \left(\frac{E_e}{m_e c^2} \right)^2 \right]$ we get

$$\dot{n}_{IC}(E_{IC}) \propto E_{IC}^{-(\alpha+1)} \int_{\xi E_{IC}}^{E_2} E_e^{-(p-2\alpha)} dE_e. \quad (9)$$

The lower bound of the integral assures with $\xi \approx 3$ scattering in the Thomson regime; the additional contribution of IC photons produced in scatterings in the KN (transition) regime has formally been absorbed into a multiplicative correction factor, which is (as discussed above) to good approximation constant as long as $\xi E_{IC} \ll E_2$. Although Eq. (9) is a crude approximation, we find that it predicts the spectral slopes satisfactorily for a wide range of p and α values provided $\alpha > 0$ (see Zdziarski 1989 who discusses in detail the applicability of the δ -functional approximation). For an electron population with spectral index p interacting with its own synchrotron radiation, i.e. $\alpha = (p-1)/2$, one obtains the well known result

$$\dot{n}_{IC}(E_{IC}) \propto E_{IC}^{-(\alpha+1)} \ln(E_2/(\xi E_{IC})), \quad (10)$$

namely that the IC spectrum has approximately the same spectral index α_{IC} as the synchrotron spectrum. A cooled electron spectrum with $p' \approx 3.2$ which interacts with the seed photon spectrum of the uncooled electrons with $\alpha \approx 0.6$ yields

$$\dot{n}_{IC}(E_{IC}) \propto E_{IC}^{-(\alpha'_{IC}+1)} \quad (11)$$

with $\alpha'_{IC} = p' - \alpha - 1 \approx 1.6$. Remarkably, Eqs. (10) and (11) show that in the case of SSC scenarios of TeV blazars changes of the electron spectral index are reflected in an approximate one to one relationship in changes of the spectral index of the produced IC component.

Based on qualitative arguments about the stability of the Mkn 501 VHE spectrum Konopelko et al. (1999b) inferred an intrinsic TeV spectrum following a pure power law with a spectral index α' near one. Given the spectral indices $p = 2.2$ and $p' = 3.2$ of the low and the high energy electrons respectively, as deduced from the X-ray spectra, Eq. (9) as well as our simulations show that it is extremely difficult to obtain such an IC spectrum in a SSC scenario since the resulting spectrum is characterized by $\alpha_{IC} = 0.6$ or $\alpha'_{IC} = 1.6$, i.e. it is too soft or too hard by ~ 0.5 units in the spectral index. Furthermore, in the scenarios discussed here, the narrow frequency range where the IC spectrum turns over ($\alpha \approx 1$) is a region where we expect more rather than less spectral variability than below and above this frequency range.

Since (i) the RXTE spectral index of April 13th approximately represents the mean value during the 1997 RXTE observations of this source and (ii) the HEGRA VHE spectra did not vary during 1997 within statistical errors we think it is justified to compare the prediction of VHE spectrum of this day with the observed 1997 time averaged VHE spectrum. The predicted VHE spectrum mainly depends on the parameter $(\delta_j/B)^{1/2}$ determining the ratio of the IC SED peak frequency and the square root of the synchrotron SED peak frequency (Eq. (7)). The Doppler factor as well as the magnetic field are poorly constrained towards higher values. These large uncertainties however have only a limited impact on the model fits. Combinations of large Doppler factors and small magnetic fields are constrained by the “mapping relation” between X-ray and VHE photon energies given in Eq. (7). Using an extremely high Doppler factor of $\delta = 100$ together with the limiting lower value of the magnetic field from Eq. (4) of 0.012 G results in a mapping of the 1 TeV to 5 TeV radiation with the 0.02 keV to 0.6 keV radiation, which seems unlikely, given the BeppoSAX, HEGRA, and CAT observations during April 1997 which showed dramatic variability above several keV and at TeV energies but almost no variability at all below 1 keV. The dashed line in Fig. 9 shows the predicted VHE spectrum for this limiting upper value of the parameter $(\delta/B)^{1/2}$. The difference between the predicted and the observed γ -ray spectrum can be interpreted as an *upper limit* on the increase of intergalactic extinction with γ -ray energy. Additionally, the magnetic field is not constrained towards higher values. The dotted line in Fig. 9 shows the effect of increasing the magnetic field from 0.037 G to a limiting value of 0.12 G (which has the same effect on the shape of the predicted IC spectrum as decreasing the Doppler factor to a value of $\delta \simeq 8$). Larger magnetic field values (more precisely, smaller values of $(\delta_j/B)^{1/2}$) are highly improbable since the resulting intrinsic VHE spectra are steeper than the observed ones in the energy range above ~ 1 TeV. Intergalactic extinction which decreases for increasing energy is rather unlikely given our present understanding and knowledge of the DEBRA density. Our upper limit on the magnetic field

11651.F10

Fig. 10. The increase of the optical depth (OD) for γ -rays due to intergalactic extinction as inferred from comparing the observed VHE Mkn 501 spectrum with the estimated emitted spectrum (from Fig. 9). The 1997 Mkn 501 time-averaged spectrum measured by HEGRA has been used and the corresponding statistical errors are shown by the vertical error bars. Additional uncertainties, arising from the assumed value for the magnetic field B and the Doppler factor δ_j are shown by the solid lines (solid points: $\delta_j = 25$, $B = 0.037$ G, upper solid line: $\delta_j = 100$, $B = 0.012$ G, lower solid line: $\delta_j = 25$, $B = 0.12$ G). The dashed line shows the prediction by Primack et al. (1999) based on the cosmological model “ Λ CDM” with cold dark matter. The dotted lines show a range of predicted optical depths from (Malkan & Stecker 1998; Stecker 1999).

is in good agreement with earlier results of Hillas (1999). Even for $B = 0.12$ G intergalactic extinction is needed to explain the data above ~ 10 TeV. Precision measurements of TeV blazar spectra in the energy region from 50 GeV to several hundred GeV, where only negligible intergalactic extinction is expected, would allow one to select the right value of $(\delta_j/B)^{1/2}$ for individual sources. Such measurements will become feasible with next generation Cherenkov telescope systems as HESS and VERITAS and with the satellite instrument GLAST. Note that it turns out that in our SSC scenarios the energy density of electrons strongly dominates over the energy density of the magnetic field for all admitted values of B and δ_j (see also Inoue & Takahara 1996).

4.3. Implications for the intergalactic absorption

The increase of the optical depth of the DEBRA above 500 GeV as inferred from comparing the model estimate of the emitted VHE spectrum with the observed one is shown in Fig. 10 (solid points). The Doppler factor and the magnetic field have been varied as before to illustrate the effect on the inferred optical depth which amounts here to about $1/2$ unit (see solid lines). As pointed out already, the method primarily gives the dependence of the optical depth with energy rather than the absolute amount of absorption. The result is consistent with a rather energy independent extinction up to energies of about 5 TeV. From 5 TeV to 20 TeV the inferred optical depth increases by 1 to 2 units. The dashed line in Fig. 10 shows the optical depth expected from the DEBRA estimates by Primack et al. (1999) based on galaxy and star formation calculations. The dotted lines in Fig. 10 show the optical depths for the “high” and “low” intensity DEBRA models of Malkan & Stecker (1998). Their results are based on an empirical approach where luminosity dependent infrared spectra of galaxies are integrated over their luminosity and redshift distributions. The figure shows that the increase in optical depth suggested by our analysis lies in the range of the theoretical expectations. The very weak energy depen-

dence of the optical depth up to ~ 10 TeV inferred from our analysis agrees well with the predictions of Primack et al.. Our calculations clearly favor the low over the high intensity DEBRA model of Malkan & Stecker. The optical depth predicted by the latter seems to increase too fast with γ -ray energy.

5. Summary

We have analyzed a data sample of simultaneous and nearly simultaneous X-ray and VHE data. The X-ray and VHE observations were performed with typical time delays below 6 h. Since X-ray and the VHE variability is found on time scales slower $1/2$ day and the spectral indices change typically by less than 0.01 units per hour, the measurements are well suited for a meaningful correlation analysis. For the first time a data base with detailed spectral and temporal information in the X-ray and VHE bands is available which covers a considerable number of distinct flares with substantial flux increases and decreases.

The strong variability signatures in both energy bands has been used to elucidate several key aspects of the emission activity. The 3-25 keV spectral index varied on a time scale of weeks from 0.7 to 1.1, indicating that the peak of the X-ray SED shifted on this time scale from the energy region above 25 keV to below several keV and vice versa. In Pian et al. (1998b) a similar shift has been reported from observations taken in April 1997 and in May 1998 and was attributed to decline of the X-ray and VHE γ -ray activity of the source from 1997 to 1998. The observations presented here show that the shift actually occurred also during the year of increased emission itself. Furthermore, the time lag between the 3 keV and the 25 keV emission is constrained to be smaller than half a day. The data shows an excellent correlation of the X-ray fluxes at $\simeq 25$ keV with the 2 TeV flux levels. The data therefore firmly establishes the correlation of the X-ray and VHE γ -ray flux levels of Mkn 501. A time lag between the X-ray and VHE γ -ray fluxes is constrained to be smaller than one day which is consistent with earlier results (Catanese et al. 1997; A99a; Aharonian et al. 1999c). We find that the spectral variation in X-rays and in the VHE γ -rays could well be of the same magnitude, but the accuracy of the VHE measurements does not allow us to draw definitive conclusions.

We interpret the data in the framework of SSC models. Using the constraints on the Doppler factor and the magnetic field of the emission volume from the variability and correlation properties in the X-ray and VHE bands we reconstruct the electron spectrum from the observed synchrotron radiation and estimate the emitted VHE spectrum. It turns out that intrinsic source properties, namely a turnover in the electron spectrum, is expected to cause a turnover of the VHE spectrum just in the TeV energy range. This result shows that the VHE measurements give,

apart from rather robust upper limits on the DEBRA density from general arguments as described e.g. by Biller et al. (1998) and A99b, reliable information about the intergalactic extinction only when the emission mechanism is understood in full detail. Within our model scenario the comparison of observed and predicted VHE spectrum gives an optical depth due to intergalactic extinction which is constant up to energies of about 5 TeV and rises by one to two units from 5 TeV to 20 TeV. Our future work will focus on a confirmation of the results presented in this paper by detailed time dependent model studies based on the code by Coppi (1992).

Acknowledgments The analysis of the RXTE data has been possible due to the High Energy Astrophysics Science Archive Research Center Online Service, provided by the NASA/Goddard Space Flight Center. We are grateful to E. Pian for providing us with the BeppoSAX data points from April, 1997. H.K. and F.A. are grateful to W. Hofmann, J. Kirk, and H. Völk for very fruitful discussions. We thank R. Tuffs for a careful reading of the text. We thank the referee A.M. Hillas for very useful suggestions.

References

- Aharonian F.A., Akhperjanian A.G., Barrio J.A., et al., 1999a, A&A 342, 69 (*A99a*)
- Aharonian F.A., Akhperjanian A.G., Barrio J.A., et al., 1999b, A&A 349, 11 (*A99b*)
- Aharonian F.A., Akhperjanian A.G., Barrio J.A., et al., 1999c, A&A 349, 29
- Bednarek W., Protheroe R.J., 1999, submitted to MNRAS
- Bhat C.L., Koul R., Kaul R.K., et al., 1997. In: Procs. Towards a Major Atmospheric Cherenkov Detector V, De Jager O.C. (ed), Kruger Park, South Africa
- Biller S.D., Buckley J., Burdett A., et al., 1998, Phys. Rev. Lett. 80, 2992
- Bloom S.D., Marscher A.P., 1993. In: AIP Conf. Proc. 280, Compton Gamma-Ray Observatory, ed. Friedlander M., Gehrels N., Macomb D.J. (New York AIP), 578
- Bradt H.V., Rothschild R.E., Swank J.H., 1993, A&AS 97, 355
- Catanese M., Bradbury S.M., Breslin A.C., et al., 1997, ApJ 487, L143
- Coppi P.S., 1992, MNRAS 258, 657
- Coppi P.S., 1997. In: Cracow Workshop on Relativistic Jets in AGNs, eds. Ostrowski M., Sikora M., Madejski G., Begelman M., Jagellonian University Press, p. 333
- Coppi P.S., Aharonian F.A., 1999, ApJ 521, L33
- Daum A., Hermann G., Heß M., et al., 1997, Astropart. Phys. 8, 1
- Dermer C.D., Schlickeiser R., 1994, ApJ 405, 945
- Djannati-Atai A., Piron F., Barrau A., et al. 1999, A&A 350, 17
- Edelson R.A., Krolik J.H., 1988, ApJ 333, 646
- Ghisellini G., Maraschi L., Dondi L., 1996, A&AS 120, 503
- Ginzburg V.L., Syrovatski S.I., 1964, The Origin of Cosmic Rays, Pergamon Press, Oxford
- Gould J., Schröder G., 1966, Phys. Rev. Lett. 16, 252
- Hayashida N., Hirasawa H., Ishikawa F., et al., 1998, ApJ 504, L71
- Hillas A.M., 1999, Astropart. Phys. 11, 27
- Inoue S., Takahara F., 1996, ApJ 463, 555
- Kataoka J., Mattox J.R., Quinn J., et al., 1999, ApJ 514, 138
- Kirk J.G., Rieger F.M., Mastichiadis A., 1998, A&A 333, 452
- Konopelko A., Hemberger M., Aharonian F.A., et al., 1999a, Astropart. Phys. 10, 275
- Konopelko A., Kirk J., Stecker F.W., Mastichiadis A., 1999b, ApJ 518, L13
- Lamer G., Wagner S.J., 1999, A&A 331, L13
- Lamer G., Brunner H., Staubert R., 1996, A&A 311, 384
- Malkan M.A., Stecker F.W., 1998, ApJ 496, 13
- Mastichiadis A., Kirk J.G., 1997, A&A 320, 19
- Nikishov A.I., 1962, Sov. JETP 14, 393
- Pian E., Vacanti G., Tagliaferri G., et al., 1998a, ApJ 492, L17
- Pian E., Palazzi E., Chiappetti L., et al., 1998b, In: Procs. of the "BL Lac Phenomenon Workshop", Ed. L. Takalo, Turku, Finland, p. 180.
- Primack J.R., Bullock J.S., Somerville R.S., Macminn D., 1999, Astropart. Phys. 11, 93
- Quinn J., Akerlof C.W., Biller S., et al., 1999, submitted to ApJ
- Sikora M., 1997, In: Proc. 4th Compton Symposium, AIP Conf. Proc., Dermer C., Strickman M., Kurfess J. (eds), 494
- Sikora M., Begelman M.C., Rees M.J., 1994, ApJ 421, 153
- Stark A.A., Gammie C.F., Wilson R.W., et al., 1992 ApJS 79, 77
- Stecker F.W., 1999, Astropart. Phys. 11, 83
- Stecker F.W., De Jager O.C., Salamon M.H., 1992, ApJ. 390, L49
- Svensson R., 1987, MNRAS 227, 403
- Takahashi T., Tashiro M., Madejski G., et al., 1996, ApJ 470, L89
- Tavecchio F., Maraschi L., Ghisellini G., 1998, ApJ 509, 608
- Ulrich M.H., Maraschi L., Urry C.M., 1997, ARA&A 35, 445
- Zdziarski A.A., 1989, ApJ 342, 1108